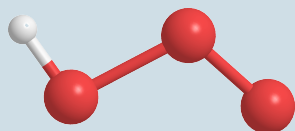


# MILLIMETER AND SUBMILLIMETER SPECTROSCOPIC STUDIES OF $\text{HO}_3$

Luyao Zou, Susanna L. Widicus Weaver

*Department of Chemistry, Emory University, Atlanta, GA, USA*



# WHY $\text{HO}_3$ ?

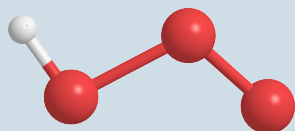
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**G**eometry

**S**pectroscopy

**B**ond Strength

**A**strochemistry



# GEOMETRY

## Central O-O bond length

### Experimental

Suma | 1.688

McCarthy | 1.684

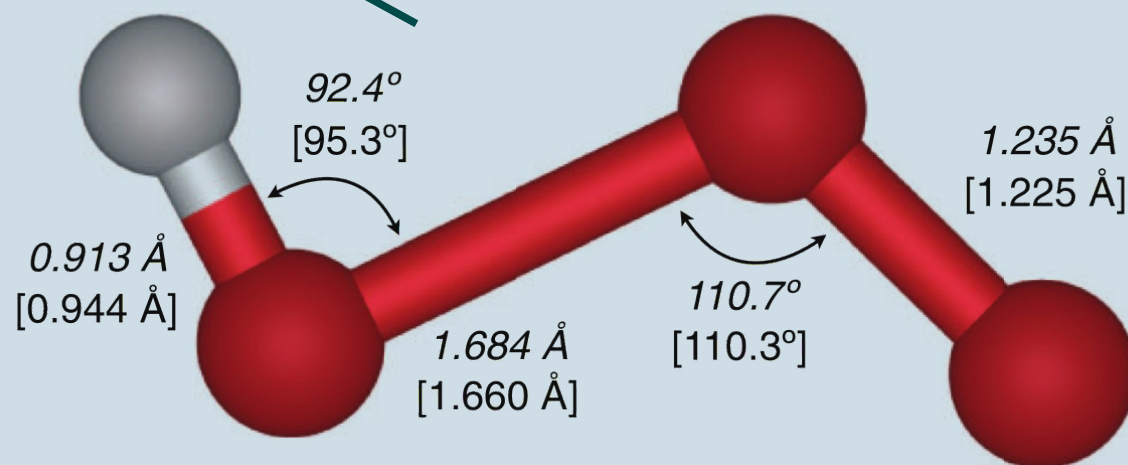
### Calculation

Hoy | 1.618

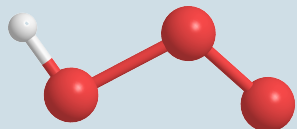
Denis | ~ 1.5

Varandas | ~ 1.42

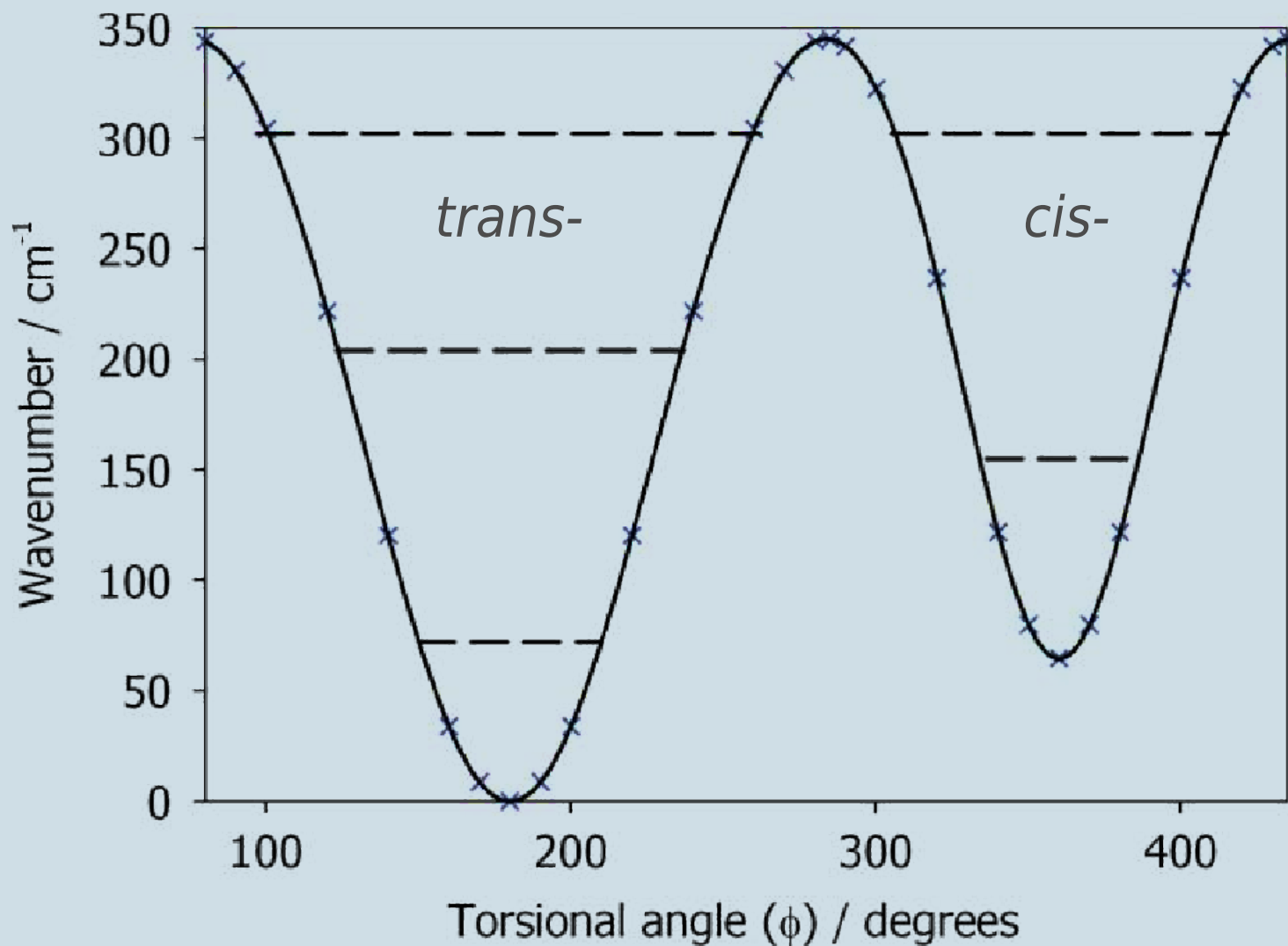
## Planar



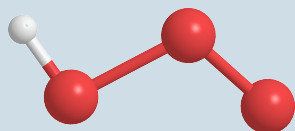
Varandas, A. J. C. et al., *Molecular Physics*, 1997  
Denis, P. A. et al., *Chem. Phys. Lett.*, 2002  
Suma, K. et al., *Science*, 2005  
McCarthy, M. C. et al., *J. Chem. Phys.*, 2012  
Hoy, E. P. et al., *J. Phys. Chem. A*, 2013



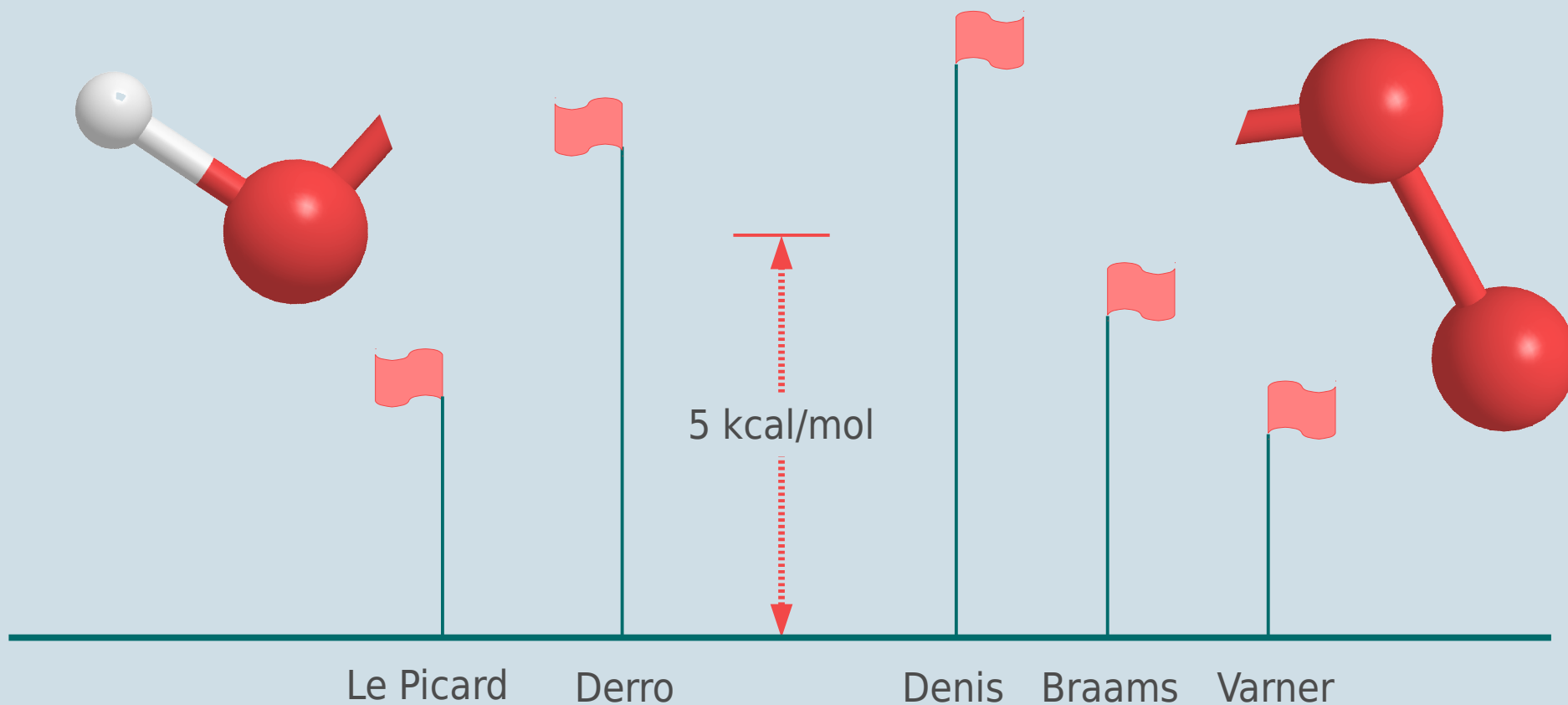
# Torsion & Isomerization



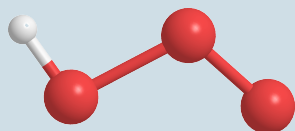




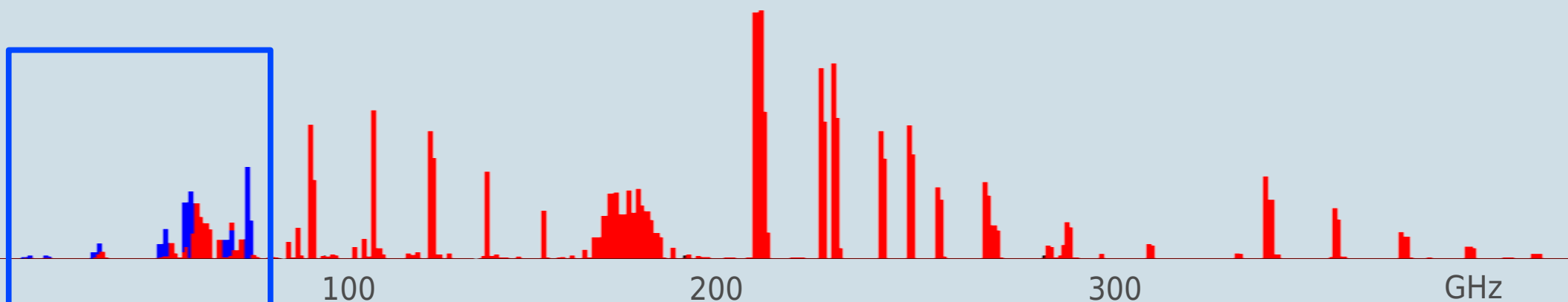
# DISSOCIATION LIMIT



Braams, B. J. et al., *Phys. Chem. Chem. Phys.*, 2008.  
Derro, E. L. et al., *J. Phys. Chem. A*, 2008  
Denis, P.A. et al., *J. Phys. Chem. A*, 2009  
Varner, E. et al., *J. Phys. Chem. A*, 2009  
Le Picard, S. D. et al., *Science*, 2010



# MEASURED TRANSITIONS



Microwave  
< 80 GHz

Millimeter  
**NONE !**

Infrared  
> 1000  $\text{cm}^{-1}$





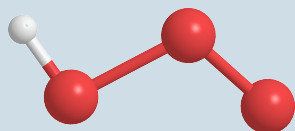
# Does **HO<sub>3</sub>** EXIST in the **Interstellar Medium** ?

Grain surface reaction in dark clouds leads to O<sub>2</sub> depletion. HO<sub>3</sub> may have a chance to form in the grain surface. However, no one has identified HO<sub>3</sub> in space.

*Orion & Horsehead Nebular  
Credit: Roberto Colombari & Federico Pelliccia,  
NASA Astronomy Picture of the Day*

*Hincelin, U. et al., Astron. & Astrophys., 2011*





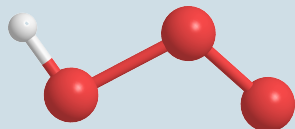
# THIS study

---

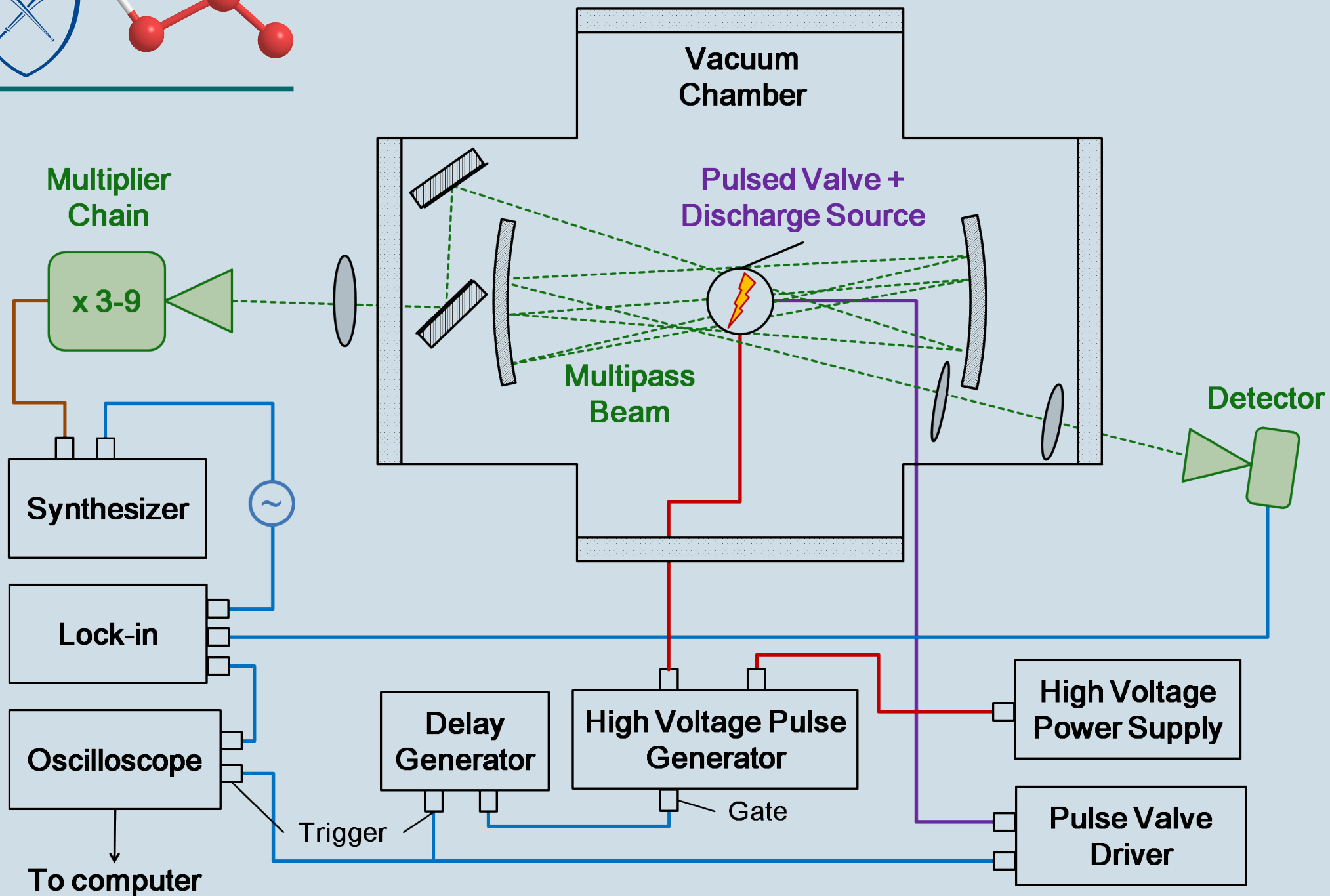
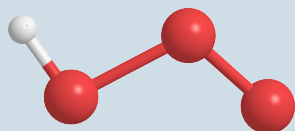
**L**INES!

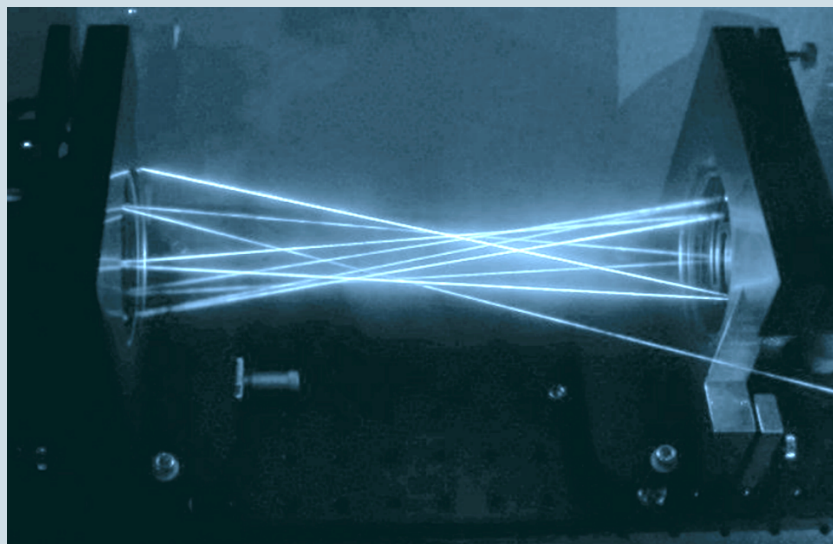
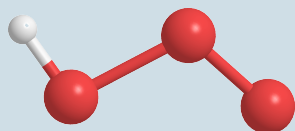
**P**arameters

**S**upport for Astrochemistry



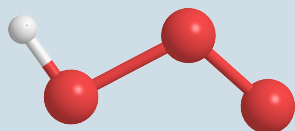
# THE **EXPERIMENT**





Direct Absorption  
Spectrometer with

**MULTIPASS**  
Optical Design



## PULSED

Discharge Source

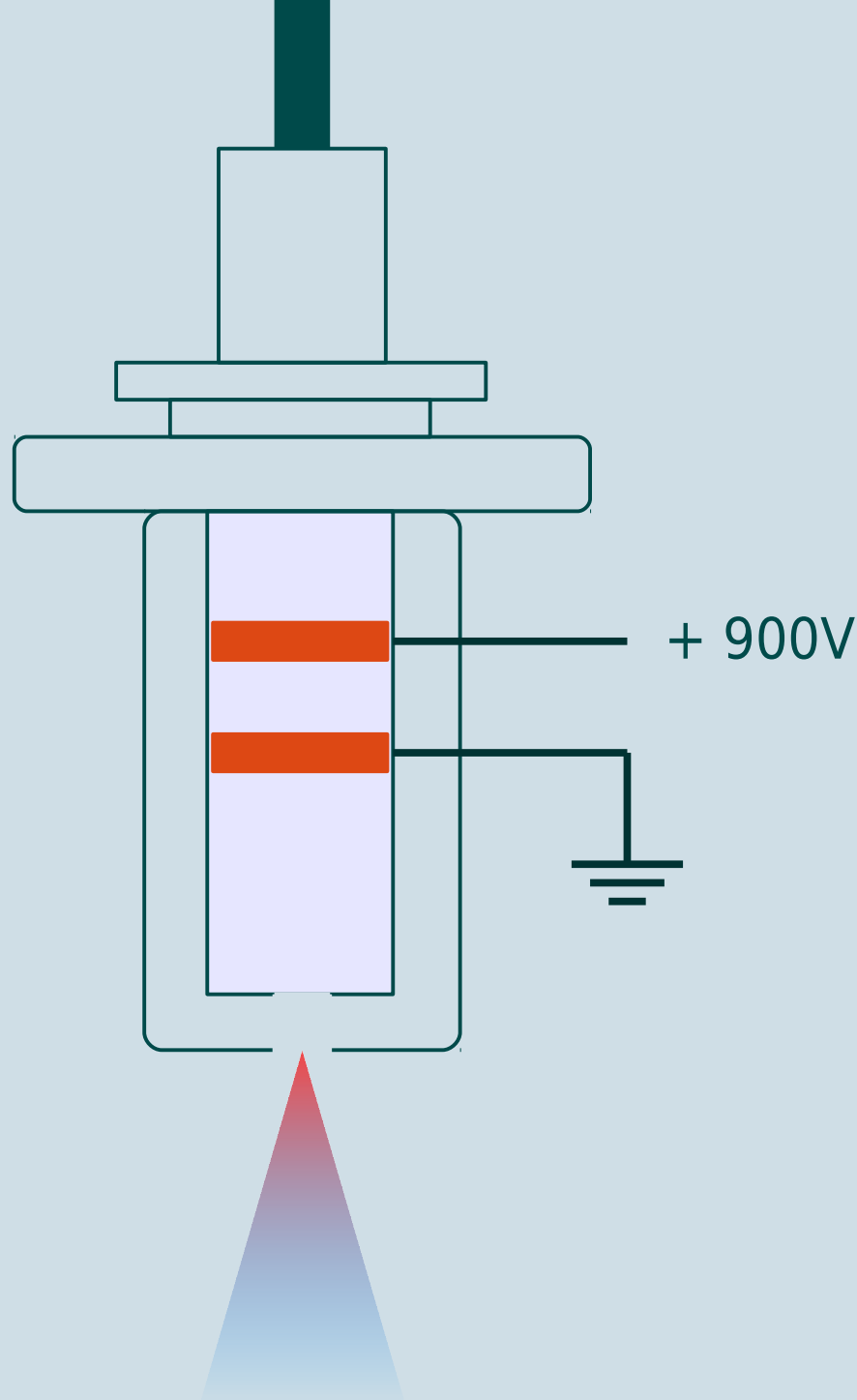
TRIGGERED by  
the pulsed valve

## SUPERSONIC

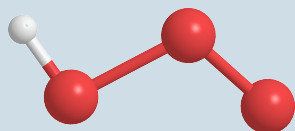
Expansion

COOLS the radical down to  
15 K rotational temperature

INCREASES line intensity

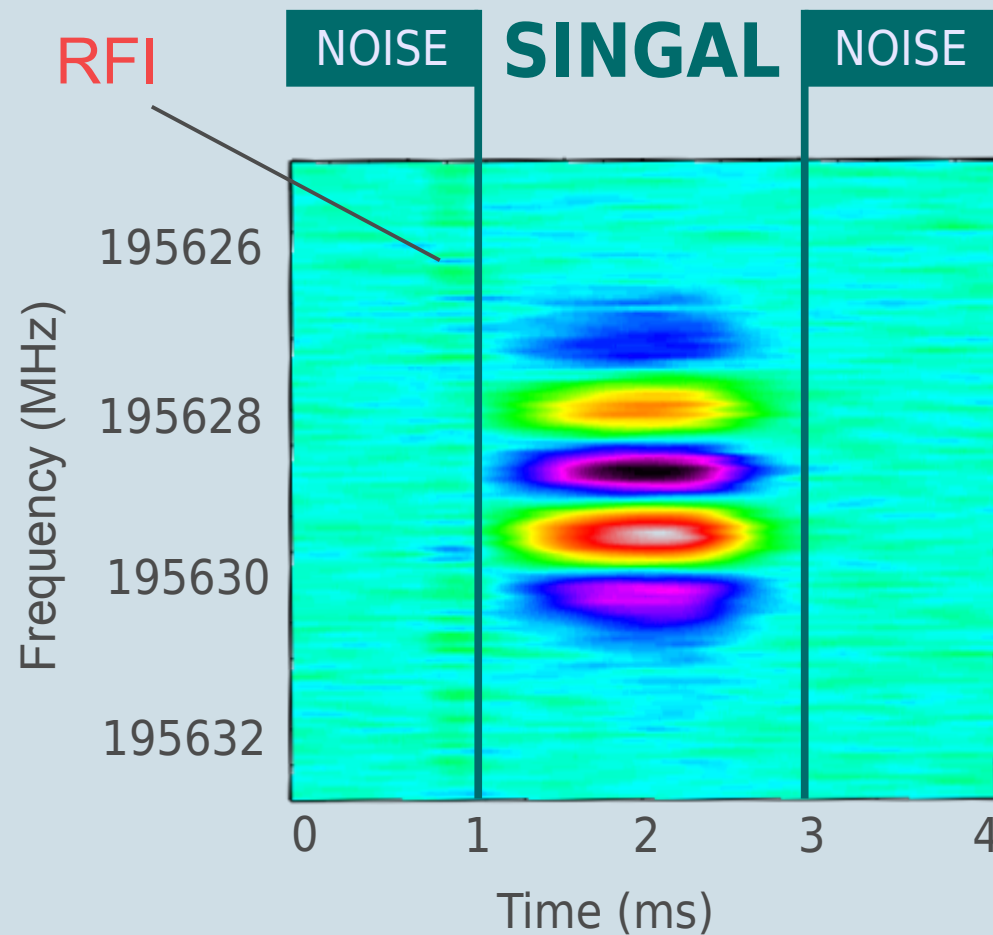


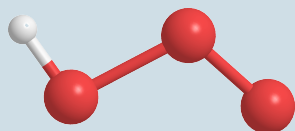




# BOXCAR Integration

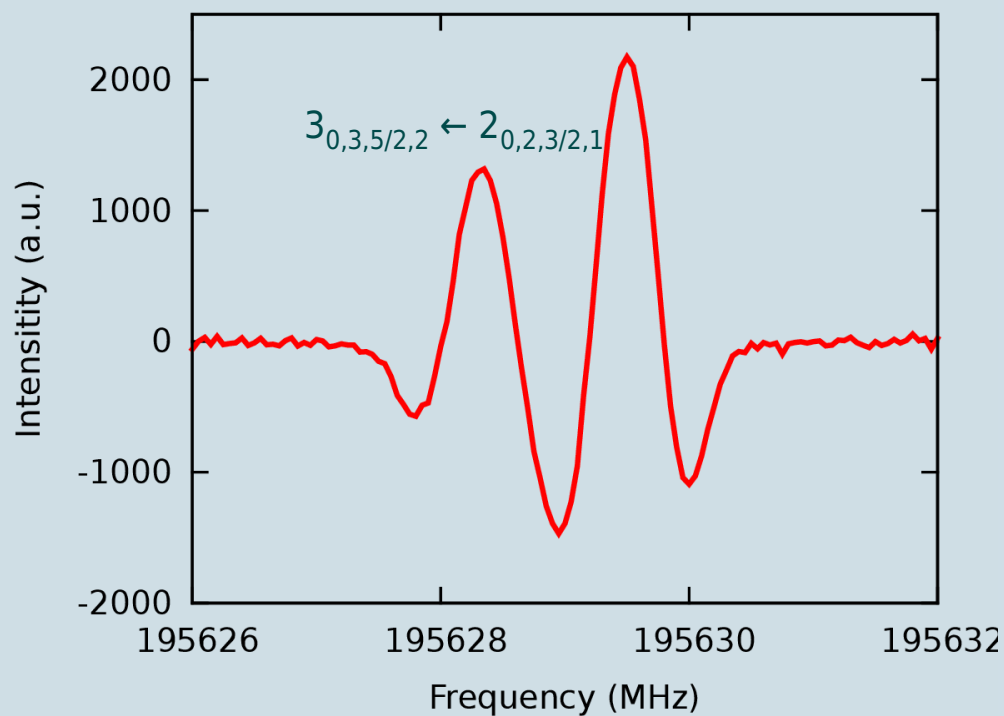
$\text{HO}_2$  lines as an **EXAMPLE**



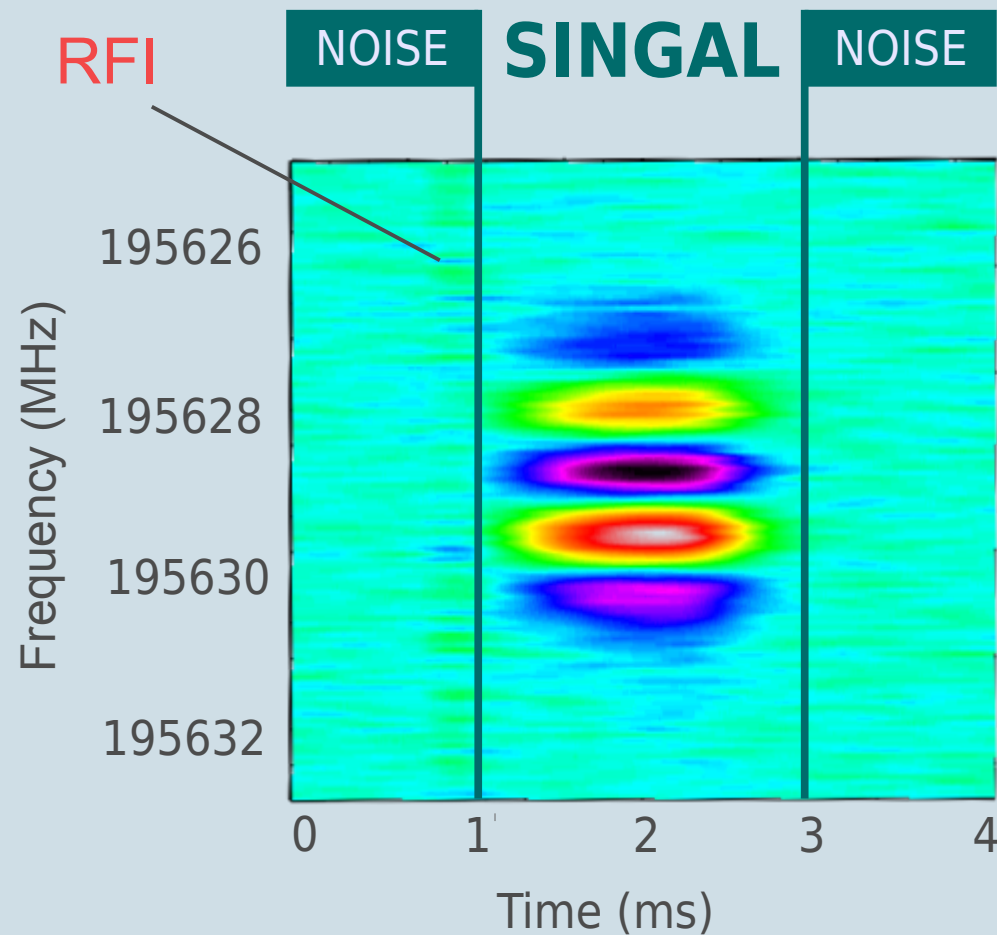


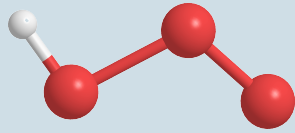
# BOXCAR Integration

$\text{HO}_2$  lines as an EXAMPLE

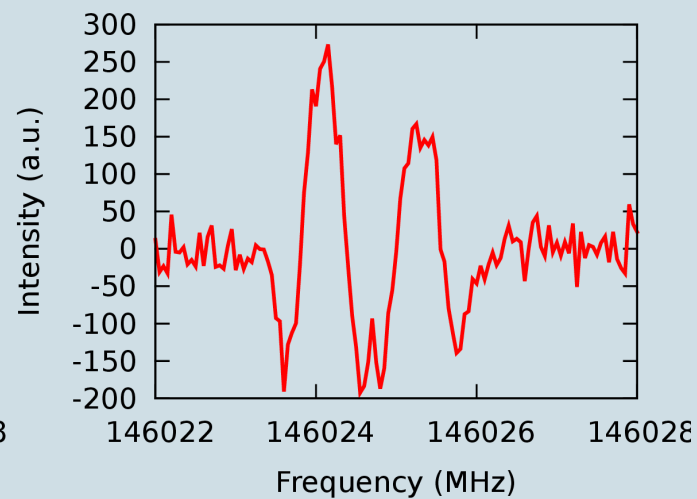
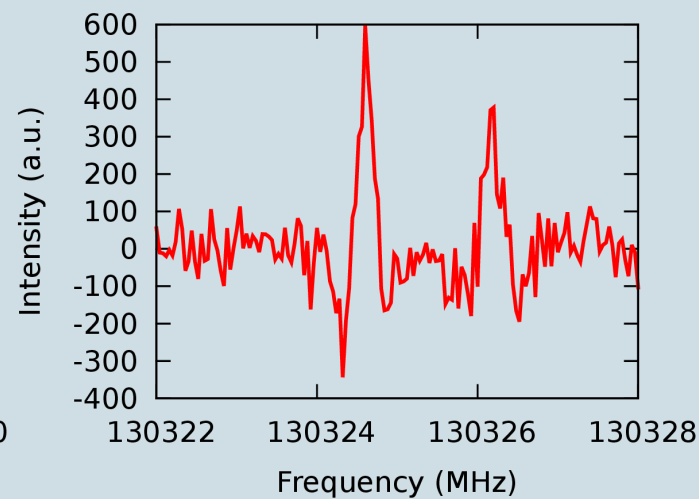
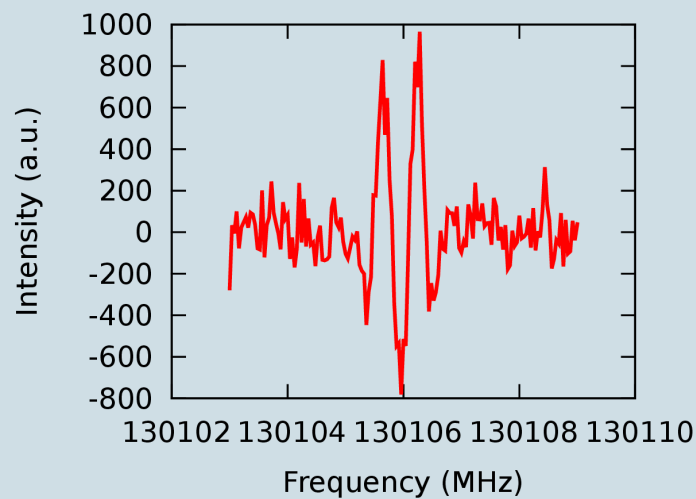
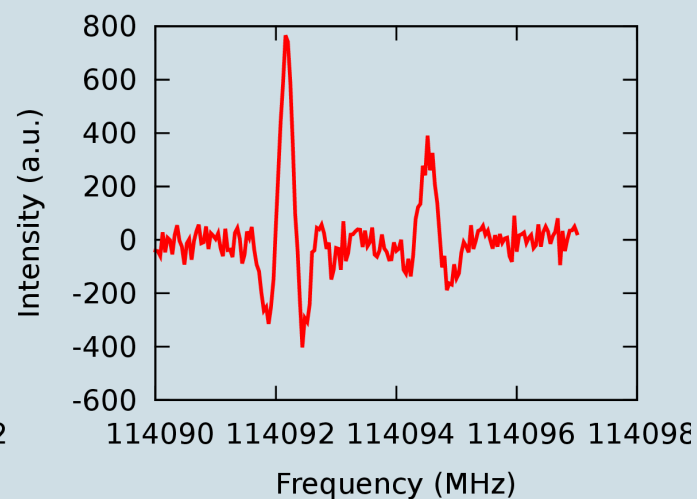
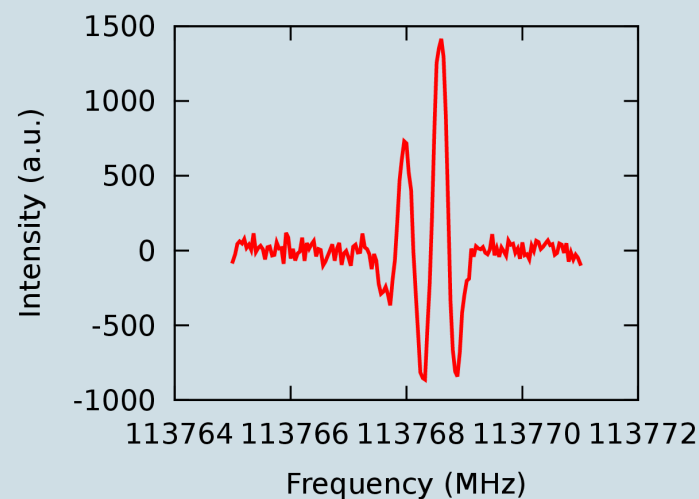
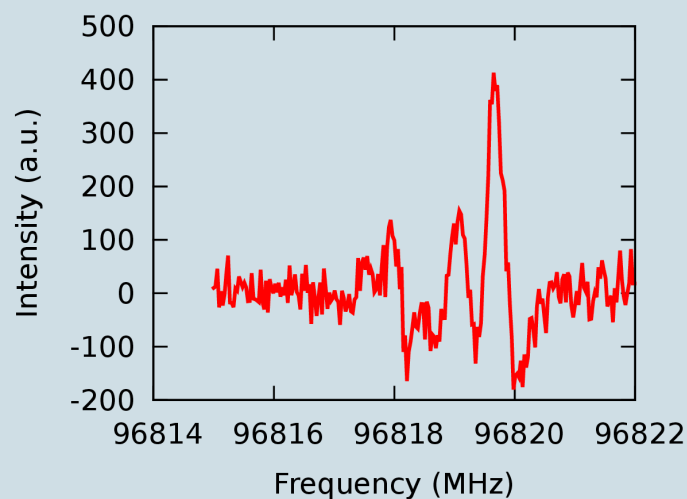
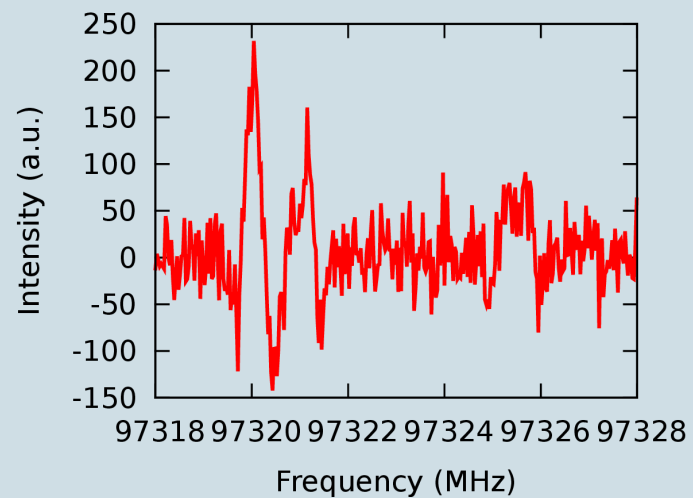
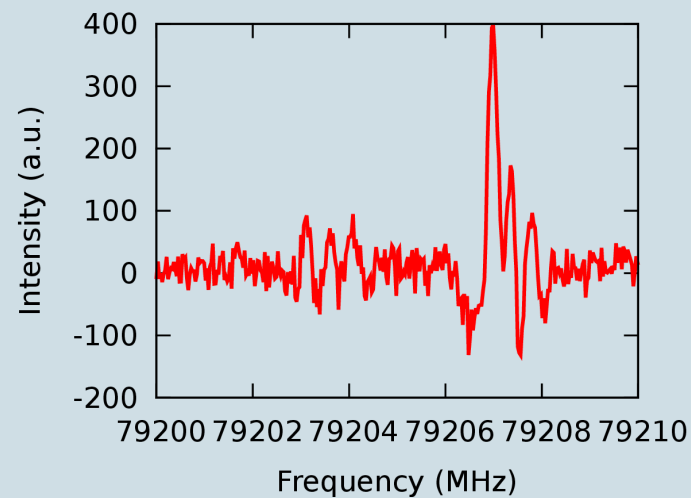
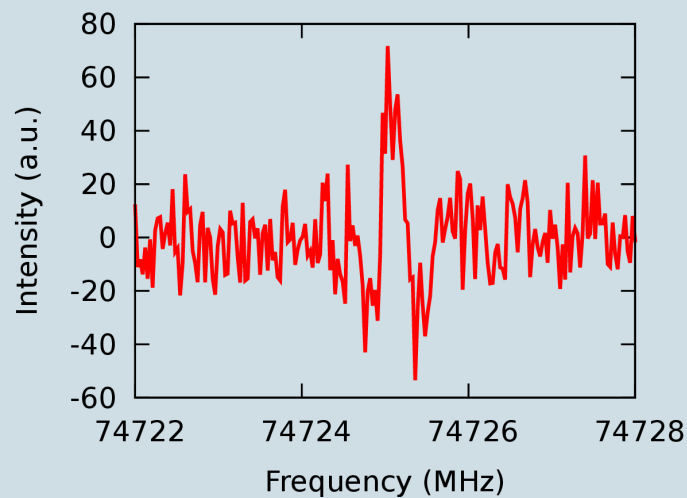


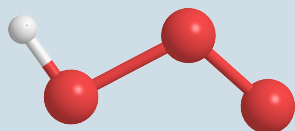
$\text{HO}_2$  transition observed readily in discharge



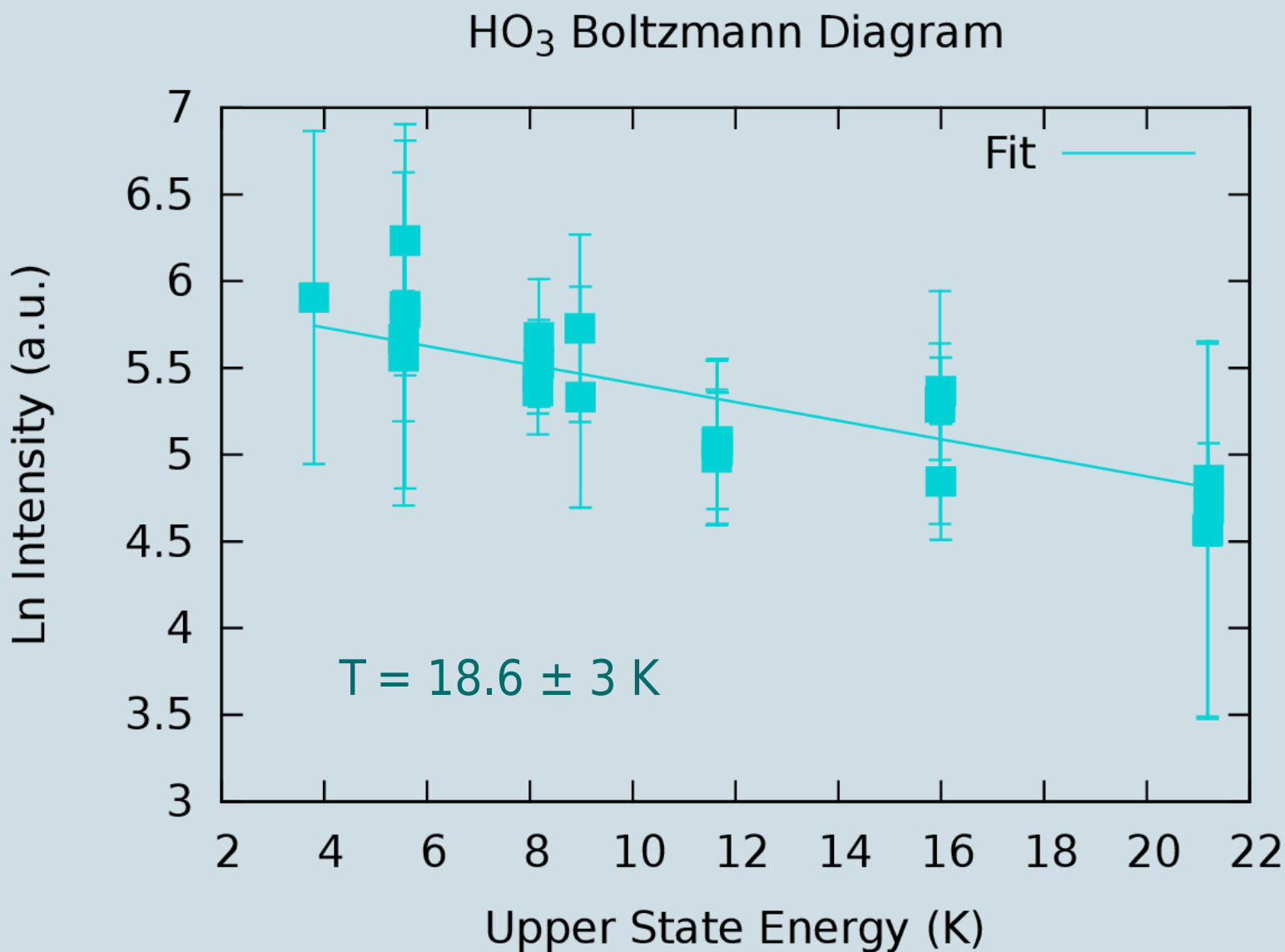


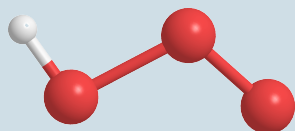
# THE $\text{HO}_3$ LINES





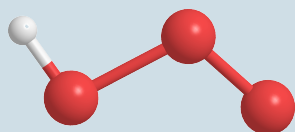
# PROBING **TEMPERATURE**



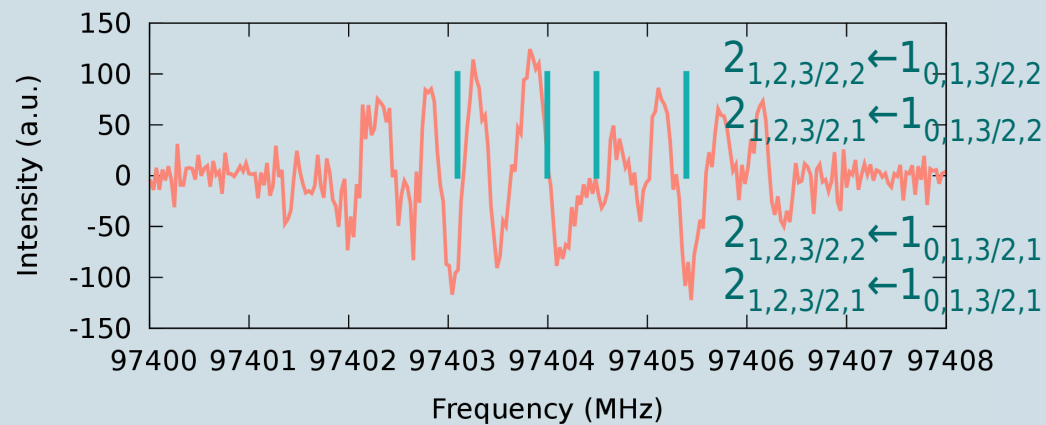
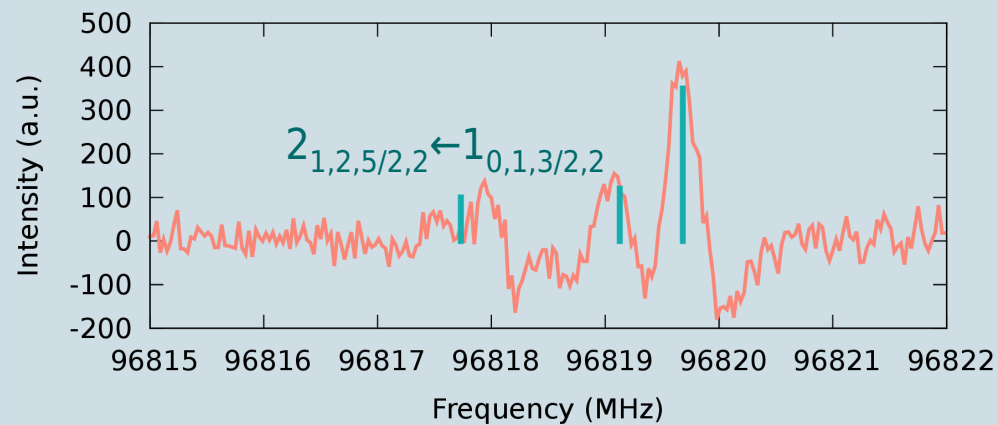
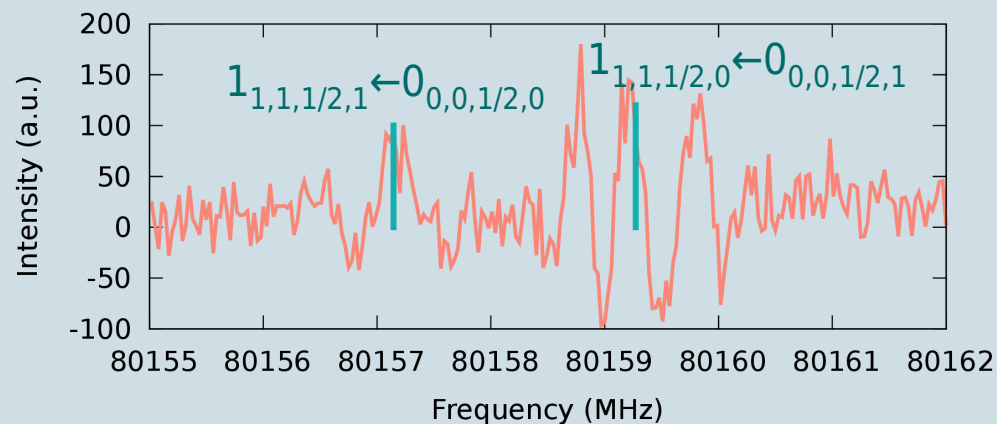
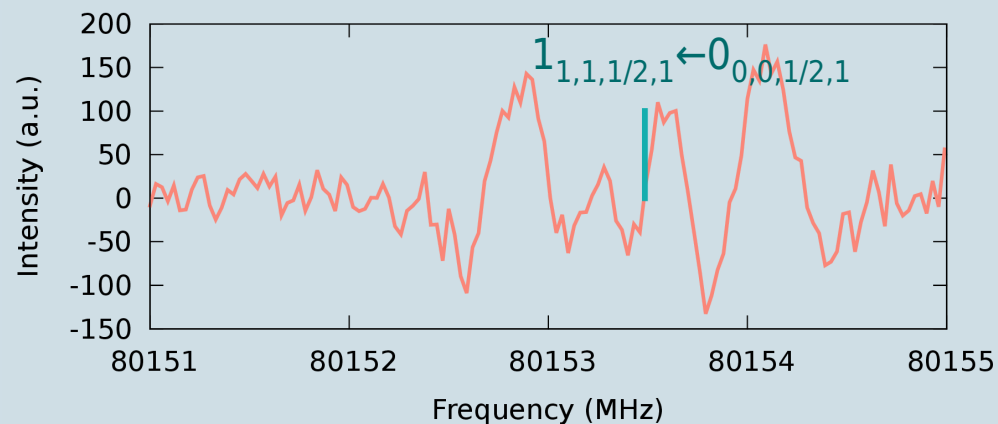
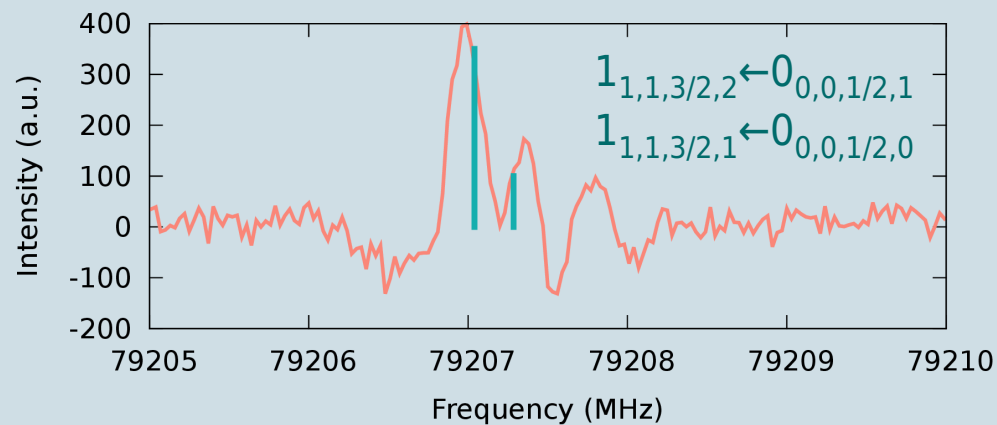
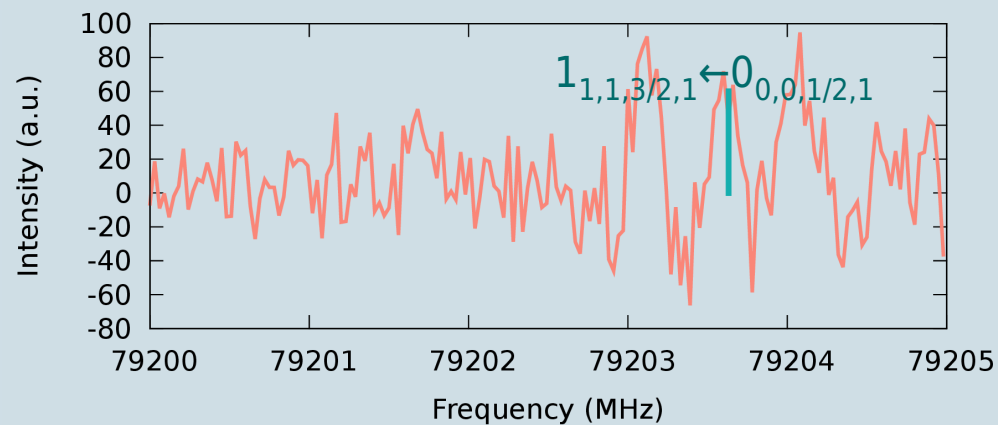


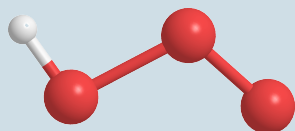
# FITTING RESULTS

Parameters	Suma et al.	This study
$A-(B+C)/2$	61409.61111(172)	61409.61043(169)
$(B+C)/2$	9368.55408(46)	9368.55403(45)
$(B-C)/4$	309.19802(39)	309.19756(34)
$\Delta_N$	0.0464608(242)	0.0464512(235)
$\Delta_{NK}$	0.15335(32)	0.153158(308)
$\delta_N$	0.006108(33)	0.0060665(281)
$\epsilon_{aa}$	-1252.5858(47)	-1252.5858(46)
$\epsilon_{bb}$	-106.25509(207)	-106.25513(205)
$\epsilon_{cc}$	-3.49541(193)	-3.49515(192)
$ \epsilon_{ab} + \epsilon_{ba} /2$	42.450(85)	42.457(84)
$\Delta_N^S$	0.000998(82)	0.0001011(80)
$\Delta_{NK}^S$	0.05441(122)	0.05449(120)
$a_F$	3.65871(243)	3.65869(243)
$1.5T_{aa}$	12.7242(71)	12.7233(71)
$0.25(T_{bb} - T_{cc})$	-1.30597(237)	-1.30628(236)



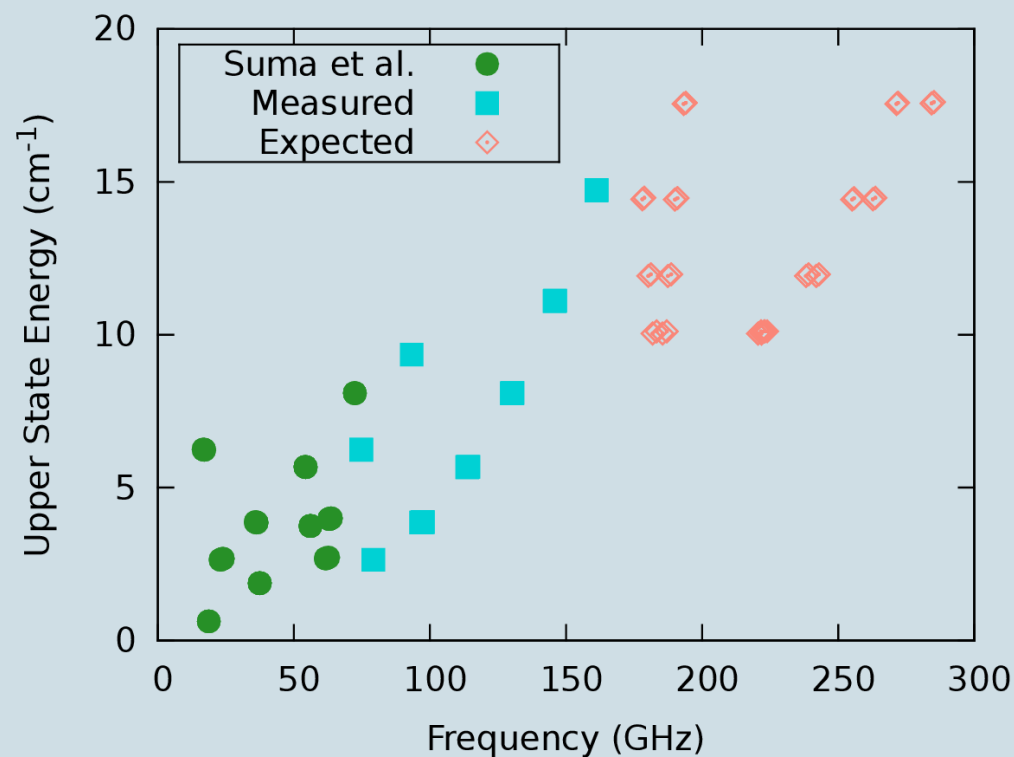
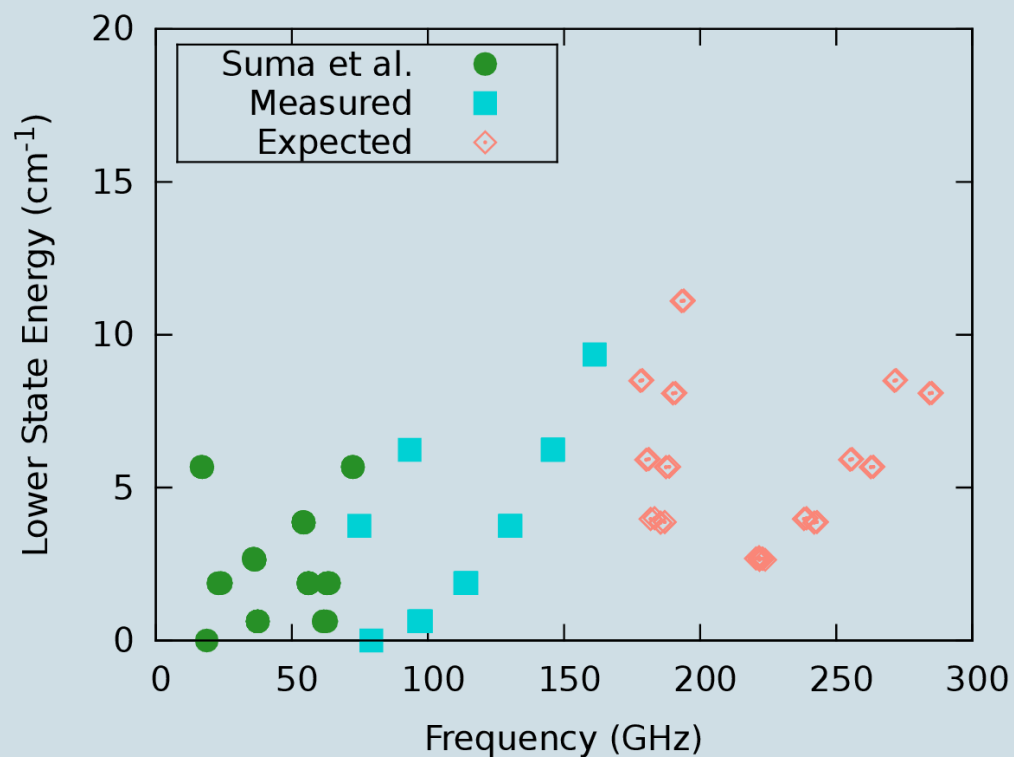
# UNEXPECTED SPLITTING



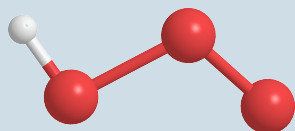


**UNABLE** TO SEE MORE LINES

Power Drop **OR** Torsion **OR** ? ? ?







## NEW LINES

>25 new lines of  $\text{HO}_3$   
are observed in the  
laboratory from 70 to  
162 GHz up to  $J=6$

## FITS

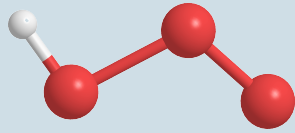
Preliminary fits are  
conducted. 15  
parameters are  
refined

## PUZZLES

Unexpected splitting  
pattern at low  
frequencies  
Unable to see more lines

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# IN SUMMARY



# What's

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# NEXT

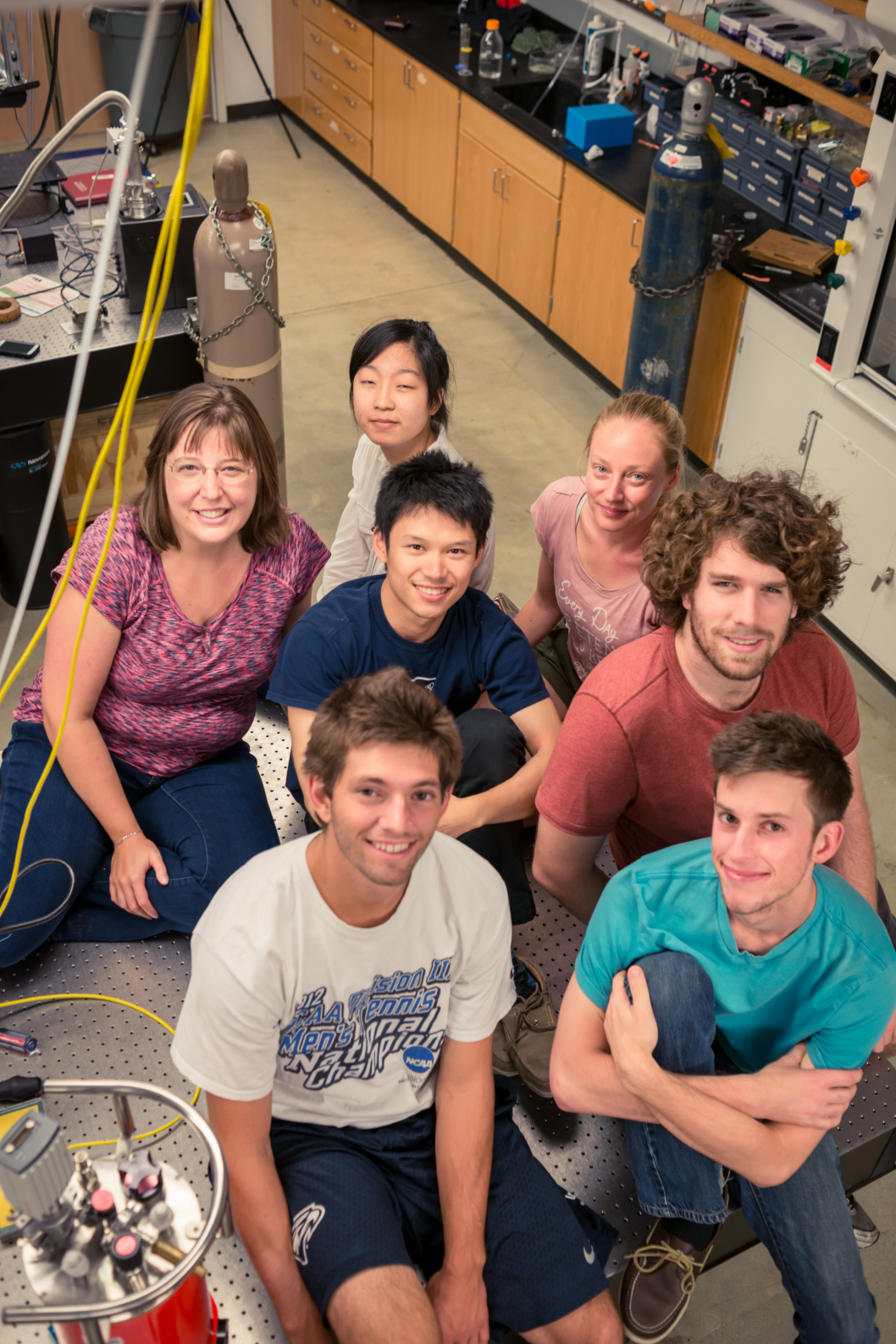
**INCREASE** spectrometer power

**REPEAT** with DO<sub>3</sub> to determine  
source of spectral splitting

**ADD** Helmholtz coils

**EXTEND** measurements to  
higher frequencies

**COMPARE** results to  
observations



# A CKNOWLEDGEMENT

The Widicus Weaver Group

Mike McCarthy

NSF and Emory University for \$\$\$

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